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X-ray monitoring of the radio and γ -ray loud Narrow-Line Seyfert 1 Galaxy PKS 2004–447

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Abstract. We present preliminary results of the X-ray analysis of *XMM-Newton* and *Swift* observations as part of a multi-wavelength monitoring campaign in 2012 of the radio-loud narrow line Seyfert 1 galaxy PKS 2004–447. The source was recently detected in γ -rays by *Fermi/LAT* among only four other galaxies of that type. The 0.5 – 10 keV X-ray spectrum is well-described by a simple absorbed powerlaw ($\Gamma \sim 1.6$). The source brightness exhibits variability on timescales of months to years with indications for spectral variability, which follows a “bluer-when-brighter” behaviour, similar to blazars.

1 Introduction

The recent detection of variable γ -ray emission in five radio-loud Narrow-Line Seyfert 1 (RL-NLS1) galaxies [1] suggests that these sources possess powerful relativistic jets, similar to blazars and “classical” radio galaxies. Narrow line Seyfert 1 galaxies (NLS1) galaxies are defined by the presence of narrow permitted optical emission lines, i.e., a strong but narrow $H\beta$ emission and a strong Fe II emission [9]. Typically Seyfert 1 galaxies are variable on timescales of months to years. On the contrary, NLS1 exhibit very strong X-ray variability in flux and spectral shape on shorter timescales down to weeks (e.g., [2]) and a strong soft excess below 2 keV. These timescales are common for radio-loud active galactic nuclei (AGN) and are associated with emission of the jet and are accompanied by flaring in other wavelengths, from radio to γ -rays. The detection of γ -ray emitting RL-NLS1 present a unique opportunity to study the similarities and differences in the emission processes and structure of radio-loud and quiet AGN, since these sources show properties of both types. The active galaxy PKS 2004–447 is a γ -ray loud NLS1 galaxy at a redshift of $z = 0.24$ [4]. First observations yield optical properties consistent with the formal NLS1 classification, (flux ratio $O\text{ [III]}/H\beta = 1.6$, $H\beta_{\text{FWHM}} = 1447 \text{ km s}^{-1}$),

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but exhibit atypical weak Fe II emission ($EW_{\text{Fe II}} \leq 10 \text{ \AA}$) [8]. The black hole mass is estimated to be $10^{6.7} M_{\odot}$ [8]. PKS 2004–447 was detected in γ -rays by the *Fermi/LAT* Telescope [1] in 2009. Following its discovery in γ -rays it has been included in the Very Long Baseline Interferometry (VLBI) TANAMI program [7] and is monitored since October 2010.

As of the time of writing, among the five known γ -NLS1 detected in 2009, only one multi-wavelength campaign has been performed on PMN J0948+0022 [1]. However, PKS 2004–447 differs dramatically from PMN J0948+0022, in terms of both physical and spectral properties, such as luminosity and black hole mass [1].

2 Observations

From May to October 2012, we conducted a multiwavelength monitoring program of PKS 2004–447 including *XMM-Newton* and *Swift* observations, providing optical/UV and X-ray data coverage, as well as quasi-simultaneous VLBI observations by TANAMI. The X-ray monitoring consists of two deep *XMM-Newton* pointings of ~ 40 ks on 2012-05-01 and 2012-10-18, which were connected by three *Swift* observations of approximately 7 ks. A follow-up observation was also performed 2013-07-07 by *Swift*. In order to study the long-term behaviour

Table 1. Results of the absorbed powerlaw fit to the X-ray data.

ObsID ^a	ObsDate ^b	Inst ^c	Γ^d	$F_{0.5-10\text{keV}}^{e,1}$	$F_{0.5-2\text{keV}}^{e,2}$	$F_{2-10\text{keV}}^{e,3}$	Stat/dof (red) ^f
0003249200[5,6]	2013-07-07	XRT	$1.49^{+0.10}_{-0.13}$	$1.07^{+0.14}_{-0.12}$	$0.29^{+0.03}_{-0.03}$	$0.79^{+0.14}_{-0.12}$	86.3/77 (1.12)
0694530201	2012-10-18	EPIC	$1.67^{+0.04}_{-0.05}$	$0.61^{+0.03}_{-0.03}$	$0.20^{+0.01}_{-0.01}$	$0.41^{+0.03}_{-0.03}$	111.9/102 (1.10)
00032492004	2012-09-30	XRT	$1.13^{+0.43}_{-0.33}$	$0.83^{+0.07}_{-0.33}$	$0.15^{+0.05}_{-0.04}$	$0.69^{+0.66}_{-0.34}$	8.3/9 (0.92)
00032492003	2012-09-01	XRT	$2.02^{+0.44}_{-0.43}$	$0.44^{+0.16}_{-0.11}$	$0.20^{+0.06}_{-0.05}$	$0.24^{+0.17}_{-0.11}$	12.9/9 (1.43)
00032492001	2012-07-03	XRT	$1.79^{+0.41}_{-0.41}$	$0.56^{+0.24}_{-0.16}$	$0.21^{+0.06}_{-0.05}$	$0.36^{+0.24}_{-0.15}$	7.1/10 (0.71)
0694530101	2012-05-01	EPIC	$1.68^{+0.06}_{-0.05}$	$0.43^{+0.03}_{-0.03}$	$0.14^{+0.01}_{-0.01}$	$0.29^{+0.03}_{-0.02}$	111.3/87 (1.27)
00091031007	2012-03-14	XRT	$1.43^{+0.36}_{-0.32}$	$0.45^{+0.25}_{-0.15}$	$0.11^{+0.04}_{-0.03}$	$0.34^{+0.26}_{-0.15}$	8.5/9 (0.94)
00091031006	2011-11-15	XRT	$1.90^{+0.40}_{-0.45}$	$0.40^{+0.19}_{-0.11}$	$0.16^{+0.04}_{-0.04}$	$0.24^{+0.19}_{-0.11}$	5.8/11 (0.52)
00091031005	2011-09-17	XRT	$1.49^{+0.23}_{-0.21}$	$1.12^{+0.25}_{-0.21}$	$0.30^{+0.06}_{-0.05}$	$0.83^{+0.26}_{-0.21}$	43.3/27 (1.6)
00091031002	2011-07-14	XRT	$1.51^{+0.45}_{-0.46}$	$0.64^{+0.35}_{-0.21}$	$0.18^{+0.06}_{-0.05}$	$0.47^{+0.36}_{-0.21}$	16.0/8 (2.0)
00091031001	2011-05-15	XRT	$1.67^{+0.41}_{-0.41}$	$0.40^{+0.18}_{-0.12}$	$0.13^{+0.05}_{-0.04}$	$0.28^{+0.18}_{-0.12}$	6.2/9 (0.68)
0200360201	2004-04-11	EPIC	$1.53^{+0.02}_{-0.03}$	$1.31^{+0.04}_{-0.04}$	$0.37^{+0.01}_{-0.01}$	$0.94^{+0.04}_{-0.04}$	222.1/258(0.86)

Notes: Best-fit parameters of the absorbed powerlaw. The absorption is fixed to its galactic absorption, $3.17 \times 10^{20} \text{ cm}^{-2}$ [6]. Uncertainties correspond to the 90% confidence limits. ^a Observation ID, [n,m] marks merged spectra. ^b Date of observation. ^c X-ray Instrument. EPIC includes both the pn and MOS 1 & 2 cameras. ^d Photon Index of the powerlaw. ^e Unabsorbed fluxes in units of $10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ for the ¹complete, ²soft and ³hard energy ranges, respectively. ^f χ^2 [C-value] per degrees of freedom (dof) and reduced value for *XMM-Newton* [*Swift*] data.

of PKS 2004–447 archival data of *XMM-Newton* and *Swift* have been included in the analysis. Here, we present first results of the data analysis of *XMM-Newton* EPIC pn/MOS and *Swift* XRT (see Table 1).

3 X-ray Analysis

Each spectrum was fitted individually with a absorbed powerlaw in the 0.5–10 keV energy range. For the galactic absorption a value of $N_{\text{H}} = 3.17 \times 10^{20} \text{ cm}^{-2}$ was assumed, based on the LAB Survey [6]. *XMM-Newton* EPIC pn/MOS spectra of each observation were rebinned to a signal-to-noise ratio (SNR) of 5, to ensure the use of χ^2 -statistics. *Swift* XRT spectra were rebinned to at least 5 spectral data counts per energy bin and evaluated with Cash statistics [3], due to their small SNR. The model yields a good fit (see Table 1). Adding freely variable excess absorption did not improve the fit significantly. Thus, to reduce the number of free parameters N_{H} was fixed to its galactic value. We tested for the existence of an Fe $K\alpha$ emission line by adding a narrow line at 6.4 keV to the *XMM-Newton* data. However, it is statistically not significant ($\Delta\chi^2 = 0.1$ for $\Delta\text{dof} = 1$) and yields an equivalent width of $\text{EW}_{6.4\text{keV}} \leq 60 \text{ eV}$ in all three observations. A soft excess cannot be confirmed in our data.

4 Results

Figure 1 shows results for the unabsorbed fluxes of all X-ray observations. A significant decrease of 50% in flux is observed between 2004 April, and the recent observations in 2011 and 2012. Furthermore, among observations in 2011–2013 there are indications for flux variations on much smaller timescales. During the monitoring (grey-shaded period), the flux increased by 50% from 2012-05-01 to 2012-10-18 and, in the following, up to $(1.07 \pm$

$0.1) \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ until July, 2013, which corresponds to a total variation of approximately $0.64 \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ in 14 months. This high flux state is similar to the one in 2004. Another possible rise in flux may be seen in the *swift* observation from 2011-09-17, where the flux increased and decreased again by approximately 50% in two months, respectively.

To investigate whether variations in the source brightness are driven by the flux or changes in the spectral shape, Fig. 2 shows the 2–10 keV flux (hard flux, hereafter) plotted against the 0.5–2 keV flux (soft flux, hereafter). A linear correlation suggests that the variations are driven by variations of the flux, instead of the spectral shape. The data are highly correlated (Spearman’s rank coefficient $\rho = 0.66$, p-value= 0.02; Pearson’s correlation coefficient $r = 0.81$, p-value= 0.002) but the large scatter and uncertainties of the *Swift* data do not allow us to draw any conclusions on the shape of the correlation and thus on the possibility of spectral variations. A linear fit to the well constrained data (*XMM-Newton* data and *Swift* XRT 0003249200[5,6]) shows a negative offset of the hard flux, while the soft flux is zero. This indicates the existence of two emission components, one where flux variations are due to changes in the normalization of the initial powerlaw, and another soft component which in comparison appears to be non-variable. If the non-variable component has a different powerlaw index, this would imply a change of the spectral index when the non-variable soft component dominates the emission. The behavior of the photon index as a function of the total X-ray flux is also given in Figure 2. The *XMM-Newton* observations in two low-state show a slightly steeper photon index ($\Gamma \sim 1.68 \pm 0.05$) compared to the 2004 *XMM-Newton* and latest *Swift* observation ($\Gamma \sim 1.53 \pm 0.02$ and $\Gamma \sim 1.49 \pm 0.10$, respectively). These variations correspond to a change of 3.0–4.8 σ and indicates a harder-when-brighter behaviour. More observations

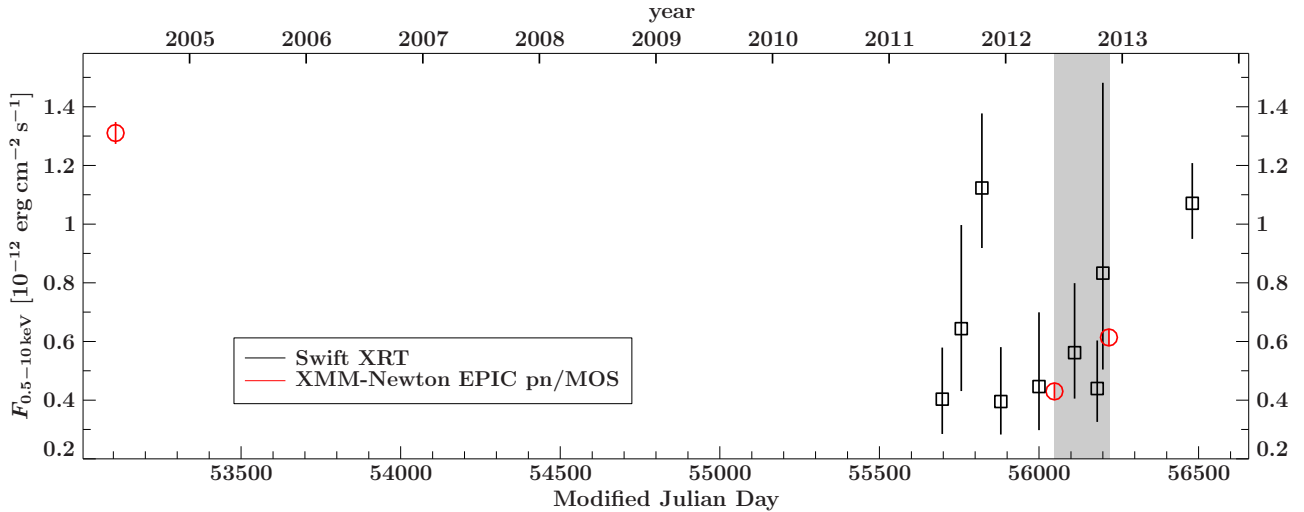


Figure 1. Unabsorbed 0.5 – 10 keV flux of PKS 2004–447 as a function of time. Red dots refer to *XMM-Newton* observations, while black squares resemble *Swift* XRT observations. The grey-shaded background marks the period of our multiwavelength campaign.

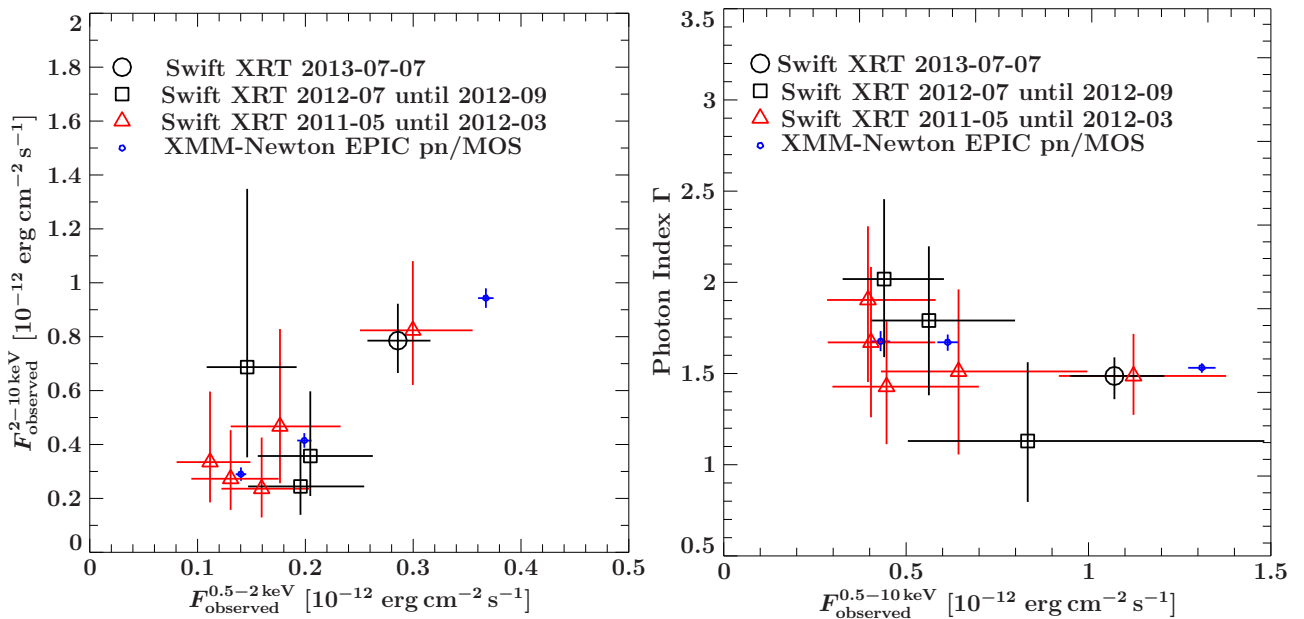


Figure 2. Hardness in the X-ray data of PKS 2004–447. *Left panel:* Unabsorbed 2–10 keV flux as a function of the 0.5–2 keV flux. *Right panel:* Photon index of the absorbed powerlaw as a function of the unabsorbed 0.5–10 keV flux.

with high Signal-to-Noise ratio spectra, however, are needed to confirm this trend and the presence of two different emission components.

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