Ultra-Heavy Cosmic Ray Analysis with CALET on the International Space Station: Established and Developing Procedures

A. W. Ficklin,¹,∗ N. Cannady,²,³,⁴ B. F. Rauch⁵ and W. Zober⁵ on behalf of the CALET Collaboration Collaboration
(a complete list of authors can be found at the end of the proceedings)

¹Department of Physics and Astronomy, Louisiana State University, 202 Nicholson Hall, Baton Rouge, Louisiana 70803, USA
²Center for Space Sciences and Technology, University of Maryland, Baltimore County, 1000 Hilltop Circle, Baltimore, Maryland 21250, USA
³Astroparticle Physics Laboratory, NASA/GSFC, Greenbelt, Maryland 20771, USA
⁴Center for Research and Exploration in Space Sciences and Technology, NASA/GSFC, Greenbelt, Maryland 20771, USA
⁵Department of Physics and McDonnell Center for the Space Sciences, Washington University, One Brookings Drive, St. Louis, Missouri 63130-4899, USA

E-mail: afickl3@lsu.edu

The CALorimetric Electron Telescope (CALET) has collected over 60 months of uninterrupted data on the flux and spectrum of the Ultra-Heavy (UH) cosmic rays from Z=30 to 40. Using the latest data provided from CALET’s UH trigger, we present a newly developed UH analysis complementary to the ongoing analysis presented at this conference by Zober et al. This work introduces a new Ultra-Heavy Analysis (UHA) dataset produced from CALET production data allowing for more streamlined analysis. We detail temporal and spatial correction algorithms using both the ⁵⁶Fe and ¹⁴Si peaks to improve charge resolution in the Z > 30 region. Additionally, this work presents a new method for removing the contributions from non-relativistic/lower-Z nuclei using trajectory tracing to compute effective cutoff rigidities in place of the previously used vertical Störmmer approximation. We show that replacing the approximated cutoffs with numerically calculated effective cutoff rigidities, calculated using the IGRF13 and T05 (Tsyganenko 05) geomagnetic field models, leads to fewer events being removed from the dataset while maintaining improved charge resolution for Z > 26. Furthermore, we introduce Tarle function peak fitting to perform charge corrections needed as a result of any quenching effects. We show the most recent CALET UH results incorporating these improvements in the analysis.

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1. Introduction

The CALorimetric Electron Telescope (CALET) has collected over 60 months of uninterrupted data on the flux and spectrum of the Ultra-Heavy (UH) cosmic rays from $Z=30$ to 40. Using the latest data provided from CALET’s UH trigger, this work presents a newly developed UH analysis complementary to the ongoing analysis by fellow CALET collaborators. We detail temporal and spatial correction algorithms using both the $^{26}$Fe and $^{14}$Si peaks to improve charge resolution in the $Z \geq 30$ region. Additionally, this work presents a new method for removing contamination from non-relativistic/lower-Z nuclei using trajectory tracing to compute effective cutoff rigidities in place of the previously used vertical Stöermer approximation. We show that replacing the approximated cutoffs with numerically calculated effective cutoff rigidities, calculated using the IGRF13 and T05 (Tsyganenko 05) geomagnetic field models, leads to fewer events being removed from the dataset while maintaining improved charge resolution for $Z > 26$. Furthermore, we introduce Tarle function peak fitting to perform charge corrections needed as a result of any quenching effects. We show the most recent CALET UH results incorporating these improvements in the analysis.

2. The CALET Instrument

CALET was launched in August 2015 to the International Space Station (ISS) for the primary purpose of directly measuring the cosmic-ray total electron spectrum up to energies of tens of TeV. The CALET calorimeter (CAL) [2], shown in Figure 1, has a depth of 30 radiation lengths ($X_0$) at normal incidence and includes three separate sub-detectors. The Charge Detector (CHD) consists of two layers of 14 plastic scintillating paddles and can measure the charge of nuclei up to $Z = 40$. The Imaging Calorimeter (IMC) contains 16 layers of 448 scintillating fibers each which provide accurate tracking for most events. Most of the shower energy is deposited in the Total Absorption Calorimeter (TASC), with 12 layers of 16 PWO logs that can contain electromagnetic showers up to TeV energies.

For ultra-heavy cosmic ray observations CALET uses a dedicated trigger mode that is active throughout the ISS orbit [1]. Standard trigger modes require events pass through the TASC, limiting the field of view, but the UH trigger removes this restriction and only requires that events pass through the CHD and top half of the IMC [1]. This raises the geometric factor from $\sim 1200\,cm^2sr$ to $\sim 4000\,cm^2sr$, and prevents UH event candidates similar to the one shown in Figure 2 from being removed from the dataset.
3. Data Analysis

3.1 Ultra-Heavy Analysis Dataset

To aid the UH data analysis a reduced high level analysis UH dataset was produced using 5 years of the CALET Level 2 Pass 4.1 dataset from October 2015 to September 2020. Included for each of the 1.8x10^8 events is the associated summary information, energy-deposit sums for each layer of the instrument, orbital parameters of the ISS at the time of observation, UH track reconstruction parameters, and individual energy deposits for detector elements on or near the calculated track. The estimated charge of each event is calculated using the event's \( dE/dx \) and the Bethe equation, which reduces to Equation 1 assuming the particle is minimum ionizing.

\[
\frac{dE}{dx} \propto z^2
\]

A selection is then made to the initial set of events requiring their reconstructed track’s incidence angle be less than 60 degrees [3]. These high incidence angle events can have their charges over estimated resulting in a reduction of charge resolution in the UH region. A charge consistency cut [3] is applied to the remaining events removing events where the estimated charges in the x and y layers of the CHD (CHDX amd CHDY) differ by greater than 1.5%. The final dataset used for this analysis includes 5x10^7 events.

3.2 Charge Corrections

3.2.1 Time/Spatial Dependence

When calculating charge using the method described above, long-term changes in temperature and PMT gain cause time dependent fluctuations that negatively impact charge resolution. To address these variations, an example of which is shown in the left side of Figure 3, events are placed in 1-day bins and fits performed to find the location of the \(^{26}\text{Fe} \) and \(^{14}\text{Si} \) peaks for each paddle of both CHDX and CHDY. These peaks are then used to calculate linear correction parameters for
each day of data. When the corrections are applied, each bin’s charge spectrum is linearly shifted so that the new $^{26}_{\text{Fe}}$ and $^{14}_{\text{Si}}$ peak locations are equal to their mean values for the full 5-year dataset.

Further variation in charge, shown in the right of Figure 3, arises from light propagation effects and non-uniformity in the CHD that both contribute to a spatial dependence in the light yield of the detector. To account for this dependence, the time-corrected data is divided into $3.2 \text{ cm} \times 3.2 \text{ cm}$ bins in both the CHDX and CHDY. For each bin, the fitted $^{26}_{\text{Fe}}$ and $^{14}_{\text{Si}}$ peak locations are again used to calculate the linear correction needed to shift each peak to its mean value for the entire CHDX and CHDY respectively.

![Figure 3: left: Location of the $^{26}_{\text{Fe}}$ peak for paddle 5 of both CHDX and CHDY as a function of time. right: Location of the $^{26}_{\text{Fe}}$ peak for each position bin of the CHDX.](image)

### 3.2.2 Quenching Correction

In organic scintillators, it is found that large deposits of energy that are spatially concentrated will exhibit non-linear saturation of the scintillator response, or quenching. This results in underestimated charge assignments that worsens as Z increases. Using a method derived by Tarle et al., it is possible to recover the true $dE/dx$ from the actual luminosity, $dL/dx$, via

$$\frac{dL}{dx} = K \frac{(1 - F_s)dE/dx}{1 + B(1 - F_s)dE/dx} + F_s \frac{dE}{dx},$$

(2)

where $F_s$ describes the fraction of electrons produced in the ionization escape between the saturated core and unsaturated halo regions of the signal, and B and K are constants. Combining Equations 1 and 2, we get

$$\frac{dL}{dx} = T(Z^2) = \frac{B_1 Z^2}{1 + B_2 Z^2} + B_3 Z^2.$$

(3)

Inverting this result then lets us solve for the un-quenched $dE/dx$ directly from our measured luminosity, $dL/dx$, using

$$Z^2 = \frac{1}{2B_2 B_3} \left[ \frac{dL}{dx} (B_2 - B_3 - B_1) + \sqrt{\left( \frac{dL}{dx} (B_2 - B_3 - B_1)^2 + 4 \frac{dL}{dx} B_2 B_3 \right)} \right],$$

(4)
where $B_1$, $B_2$, and $B_3$ are fit constants. To calculate the necessary Tarle function parameters, we performed a multi-peak fit for all nuclei from $^6$He to $^{28}$Ni and fit these directly with the Tarle function. The result of applying this quenching correction is shown in Figure 4.

![Position/Time Corrected Dequenched UH Spectra](image)

**Figure 4:** Quenching corrected CHDX charge spectra for time and position corrected UH events.

### 3.3 Charge Resolution Improvements

#### 3.3.1 Stoermer Cutoff Rigidity

To achieve adequate charge resolution in the UH ($Z \geq 30$) region, a method was required for removing the primary source of contamination, non-relativistic and lower $Z$ nuclei that are below minimum ionizing and consequently have significantly increased $dE/dx$. Given CALET’s orbit is deep within the Earth’s geomagnetic field and experiences a wide range of cutoff rigidities, implementing a cut using geomagnetic cutoff rigidity was the chosen method. Non-relativistic nuclei below minimum ionization have access to the instrument through a smaller, lower range of rigidities than higher energy nuclei, so removing times where the cutoff rigidity is in this range limits their inclusion in the dataset.

Initially, Stoermer’s vertical cutoff rigidity approximation [4], Equation 5, was used for determining cutoffs where $\Psi$ is magnetic latitude.
\[ R_c = 14.9 \times \cos^4(\Psi), \] (5)

Figure 5 shows the results of implementing a cut using this approximated cutoff, showing cut thresholds of 3 GV and 4 GV. Although there is visible improvement in the charge resolution, the loss of statistics in the UH region limits quickly becomes prohibitive to the analysis.

To refine the calculation of the cutoff rigidities, the vertical Stoermer approximation was replaced with its trajectory dependent equivalent [4]. Instead of assuming normal incidence, Equation 6 now takes into account the incident particle’s zenith, \( \phi \), and azimuth, \( \theta \), angles, as well as its distance from the magnetic dipole, \( r \), when deriving the cutoff. A comparison between the computed cutoff rigidities is shown in Figure 6, and shows significant divergence between the vertical and trajectory dependent values.

\[ R_c = \frac{60}{r^2} \times \left( 1 - \sqrt{1 - \cos(\Psi)\cos^3(\theta)} \right)^2 \] (6)

This divergence though appears primarily at times of high cutoff rigidity, and does not affect events observed over much of the orbit.

3.3.2 Numerical Method for Calculating Cutoff

To improve on the Stoermer approximation, a trajectory tracing method developed by Smart and Shea [5] was implemented to numerically solve for effective cutoff rigidities on an event by event basis. For each trajectory, the method uses an adaptive step size Runge-Kutta method to integrate the equations of motion for a charged particle in a magnetic field, stepping through the trajectory until it is deemed successful or forbidden. This is repeated for a range of rigidity values for each event, resulting in a spectrum of allowed and forbidden rigidities. An example of this, shown in Figure 7, shows all allowed rigidities in white and forbidden rigidities in blue. Notable points are the upper cutoff rigidity where the first forbidden trajectory occurs, and the lower rigidity where the last allowed rigidity occurs. The effective cutoff rigidity is calculated using

\[ R_E = R_L + \left( R_U - R_L \right) \times \frac{N_{Forbidden}}{N_{Total}}, \] (7)
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Figure 7: Spectrum of allowed and forbidden trajectories for an example $^{56}$Fe event trajectory, with allowed trajectories being white space, and forbidden trajectories being blue space.

where $R_L$ is the lower cutoff, $R_U$ is the upper cutoff, $N_{forbidden}$ is the number of forbidden trajectories after $R_U$, and $N_{total}$ is the total number of trajectories between $R_U$ and $R_L$. This effective rigidity is the value used for the calculated cutoff rigidity, as it takes into account the effect of the penumbra (allowed trajectories interspersed amongst forbidden trajectories) and is a more appropriate value when compared to the upper or lower cutoff. The effective cutoffs were calculated for all events in the UH dataset using the IGRF13 and Tsyganenko 05 internal/external geomagnetic field models. The resulting UH spectrum, Figure 8, incorporating all the corrections discussed above, and using the effective cutoff rigidity cut, shows a clear improvement in charge resolution in the UH region.

Figure 8: UH charge spectrum with effective cutoff rigidities $> 4$ GV.
4. Conclusions

The analysis improvements described in this work have provided results consistent with previous CALET UH results, producing a UH spectrum with resolved peaks for evenly charged nuclei up to Z=38. Using numerically computed effective cutoff rigidities to limit contamination in the UH dataset has improved charge resolution in the UH region. Further refinements to the analysis, including analysis presented by Zober et al. [6], along with increased statistics from the extended CALET mission will continue to improve CALET UH results going forward.

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References


Full Authors List: CALET Collaboration Collaboration


1Department of Physics, University of Florence, Via Sansone, 1, 50019 Sesto, Firentino, Italy. 2INFN Sezione di Firenze, Via Sansone, 1, 50019 Sesto, Firentino, Italy. 3Waseda Research Institute for Science and Engineering, Waseda University, 17 Kikuxicho, Shinjuku, Tokyo 162-0044, Japan. 4JEM Utilization Center, Human Spaceflight Technology Directorate, Japan Aerospace Exploration Agency, 2-1-1 Sengen, Tsukuba, Ibaraki 305-8505, Japan. 5Institute for Cosmic Ray Research, The University of Tokyo, 5-1-5 Kashiwa-no-Ha, Kashiwa, Chiba 277-8582, Japan. 6Department of Physical Sciences, Earth and Environment, University of Siena, via Roma 56, 53100 Siena, Italy. 7INFN Sezione di Pisa, Polo Fibonacci, Largo B. Pontecorvo, 3, 56127 Pisa, Italy. 8Department of Physics and McDonnell Center for the Space Sciences, Washington University, One Brookings Drive, St. Louis, Missouri 63130-4899, USA. 9Heliophysical Physics Laboratory, NASA/GSFC, Greenbelt, Maryland 20771, USA. 10Department of Physics, Catholic University of America, Washington, DC 20064, USA. 11Center for Space Sciences and Technology, University of Maryland, Baltimore County, 1000 Hilltop Circle, Baltimore, Maryland 21250, USA. 12Astroparticle Physics Laboratory, NASA/GSFC, Greenbelt, Maryland 20771, USA. 13Center for Research and Exploration in Space Sciences and Technology, NASA/GSFC, Greenbelt, Maryland 20771, USA. 14Institute of Applied Physics (IFAC), National Research Council (CNR), Via Madison del Piano, 10, 50019 Sesto, Firentino, Italy. 15Department of Physics and Astronomy, Louisiana State University, 202 Nicholson Hall, Baton Rouge, Louisiana 70803, USA. 16Department of Physics and Astronomy, University of Padova, Via Marzolo, 8, 35131 Padova, Italy. 17INFN Sezione di Padova, Via Marzolo, 8, 35131 Padova, Italy. 18Institute of Space and Astronautical Science, Japan Aerospace Exploration Agency, 3-1-1 Yoshinodai, Chuo, Sagamihara, Kanagawa 252-5210, Japan. 19Kanagawa University, 3-27-1 Rokkakubashi, Kanagawa, Yokohama, Kanagawa 221-8686, Japan. 20Faculty of Science and Technology, Graduate School of Science and Technology, Hiroshima University, 3, Bunkyo, Hiroshiki, Aomori 036-8561, Japan. 21Yukawa Institute for Theoretical Physics, Kyoto University, Kitashirakawa Oiwakecho, Sakyo, Kyoto 606-8502, Japan. 22Department of Electronic Information Systems, Shibaura Institute of Technology, 307 Fukasaku, Minuma, Saitama 337-8570, Japan. 23School of Advanced Science and Engineering, Waseda University, 3-4-1 Okubo, Shinjuku, Tokyo 169-8555, Japan. 24National Institute of Polar Research, 3-10, Midori-cho, Tachikawa, Tokyo 190-8518, Japan. 25Faculty of Engineering, Division of Intelligent Systems Engineering, Yokohama National University, 79-5 Tokiwadai, Hodogaya, Yokohama 240-8501, Japan. 26Faculty of Science, Shinshu University, 3-1-1 Asahi, Matsumoto, Nagano 390-8621, Japan. 27Hakubi Center, Kyoto University, Yoshida Honmachi, Sakyo-ku, Kyoto 606-8501, Japan. 28Department of Astronomy, Graduate School of Science, Kyoto University, Kitashirakawa Oiwake-cho, Sakyo-ku, Kyoto 606-8502, Japan. 29Institute of Particle and Nuclear Studies, High Energy Accelerator Research Organization, 1-1 Oho, Tsukuba, Ibaraki 305-0801, Japan. 30University of Pisa, Polo Fibonacci, Largo B. Pontecorvo, 3, 56127 Pisa, Italy. 31Astroparticle Physics Laboratory, NASA/GSFC, Greenbelt, Maryland 20771, USA. 32Department of Electrical and Electronic Systems Engineering, National Institute of Technology, Ibaraki College, 866 Nakane, Hitachinaka, Ibaraki 312-8508, Japan. 33Department of Physics of University of Maryland, College Park, Maryland 20742, USA. 34Department of Physical Sciences, College of Science and Engineering, Ritsumeikan University, Shiga 525-8577, Japan. 35Faculty of Science and Engineering, Global Center for Science and Engineering, Waseda University, 3-4-1 Okubo, Shinjuku, Tokyo 169-8555, Japan. 36Department of Physics and Astronomy, University of Denver, Physics Building, Room 211, 2112 East Wesley Avenue, Denver, Colorado 80208-6900, USA. 37Quantum ICT Advanced Development Center, National Institute of Information and Communications Technology, 4-2-1 Nukui-Kita-machi, Koganei, Tokyo 184-8795, Japan. 38College of Science and Engineering, Department of Physics and Mathematics, Aoyama Gakuin University, 5-10-1 Fuchinobe, Chuo, Sagamihara, Kanagawa 252-5258, Japan. 39College of Industrial Technology, Nihon University, 1-2-1 Izumi, Narashino, Chiba 275-8575, Japan. 40RIKEN, 2-1 Hirosawa, Wako, Saitama 351-0198, Japan. 41Division of Mathematics and Physics, Graduate School of Science, Osaka City University, 3-3-138 Sugimoto, Sumiyoshi, Osaka 558-8585, Japan. 42National Institutes for Quantum and Radiation Science and Technology, 4-9-1 Anagawa, Inage, Chiba 263-8555, Japan. 43Nagoya University, Furo, Chikusa, Nagoya 464-8601, Japan. 44College of Science, Ibaraki University, 2-1-1 Bunkyo, Mito, Ibaraki 310-8512, Japan.